

Asymmetric emission of the [OIII] λ 5007 profile in narrow-line Seyfert 1 galaxies [★]

E. O. Schmidt^{1,★★}, G. A. Oio¹, D. Ferreira^{1,2}, L. Vega^{1,2}, and W. Weidmann^{2,3}

¹ Instituto de Astronomía Teórica y Experimental (IATE), Universidad Nacional de Córdoba, CONICET, Observatorio Astronómico de Córdoba, Laprida 854, X5000BGR, Córdoba, Argentina.
e-mail: eduschmidt@oac.unc.edu.ar

² Observatorio Astronómico de Córdoba, Laprida 854, X5000BGR, Córdoba, Argentina.

³ Consejo de Investigaciones Científicas y Técnicas (CONICET) de la República Argentina.

February 2, 2018

ABSTRACT

Context. Many active galactic nuclei (AGN) and particularly narrow-line Seyfert 1 (NLS1) galaxies, usually exhibit blueshifts and blue wings in several emission lines, which are mainly associated with outflows and strong winds. Although there is no clear evidence of the physical origin of the [OIII] λ 5007 blue wings, it has been suggested that they could be emitted from the inner parts of the narrow-line region (NLR).

Aims. In order to study the radial velocity difference between the narrow component of H β and the core component of [OIII] λ 5007 and the asymmetric emission of this forbidden line, we investigate a sample of NLS1 galaxies. One of the aims of this paper is to analyze the blue wings of the [OIII] λ 5007 profiles and their relation with the central engine.

Methods. We have obtained and studied medium-resolution spectra (190 km s⁻¹ FWHM at H β) of a sample of 28 NLS1 galaxies in the optical range 4300 - 5200 Å. We performed Gaussian decomposition to the H β and [OIII] λ 4959,5007 emission profiles in order to study the distinct components of these lines.

Results. A new blue outlier galaxy is found, in which the center of the core component of [OIII] is blueshifted by 405 km s⁻¹ relative to the center of the narrow component of H β line. We confirmed a previously known correlation between the blueshift and the full width half maximum (FWHM) of the core component of [OIII] λ 5007 line. We also corroborated the correlation between the latter and the velocity of the centroid of the blue wing. On the other hand, by studying the radial velocity difference between the blue end of the asymmetric emission and the centroid of the core component of [OIII], we found a correlation between it and the central black hole mass and, therefore, with the luminosity of the broad component of H β . Finally, we found a moderate correlation between the luminosity of the [OIII] blue wing and the black hole mass.

Conclusions. These results suggest that the asymmetric emission of the [OIII] lines is related to the central engine, not only through the black hole but also in the intensity of the AGN, which is in agreement with previous results.

Key words. Galaxies: active – nuclei – Seyfert – kinematics and dynamics

1. Introduction

Narrow-line Seyfert 1 (NLS1) galaxies are a subclass of AGN defined according to the width of the broad component of their optical Balmer emission lines and their relative weakness of the [OIII] λ 5007 emission with full width at half maximum (FWHM) < 2000 km s⁻¹ and [OIII] λ 5007/H β ≤ 3 (Osterbrock & Pogge 1985; Goodrich 1989). Classifications based on optical spectroscopy suggest that NLS1 represent about 15% of the whole population of Seyfert 1 galaxies (Williams et al. 2002).

Related to the nuclear kinematics, the presence of blue wings and blueshifts in several high and medium degree of ionization emission lines have been reported in different AGNs and were associated with winds and outflows (e.g., Boroson & Oke 1987; Marziani et al. 2003a). These outflows can provide energy and

mass to the interstellar medium of the host galaxy (Colbert et al. 1996; Churazov et al. 2001; Moll et al. 2007).

There is observational evidence of these blueshifts in all wavelengths from the X-ray band (e.g., Chelouche & Netzer 2005; Krongold et al. 2007) to UV (e.g., Crenshaw & Kraemer 2007; Sulentic et al. 2007), optical (e.g., Das et al. 2005), IR (e.g., Rodríguez-Ardila et al. 2006), and radio (e.g., Gallimore et al. 2006).

Zamanov et al. (2002) studied a sample of 216 type 1 AGNs and found 7 of these sources with [OIII] λ 5007 blueshifts < -250 km s⁻¹ (called blue outliers). Likewise, objects with similar characteristics were analyzed by several authors (e.g., Grupe et al. 2001; Grupe & Leighly 2002; Marziani et al. 2003b; Aoki et al. 2005). Recently, Zakamska et al. (2016) detected quasars at $z \sim 2.5$ with blueshifts of ~ -1500 km s⁻¹.

Considering high ionization emission lines such as [Fe VII], [Fe X], and [Fe XI], Rodríguez-Ardila et al. (2006) analyzed blueshifts in Seyfert 1 galaxies. They found that the size of the emitting region varies with the ionization potential. In this scenario, the higher the ionization potential, the more compact the emitting region becomes. This suggests that nuclear photoion-

[★] The reduced spectra as FITS files are only available in electronic form at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsweb.u-strasbg.fr/cgi-bin/qcat?J/A+A/>

^{★★} Visiting Astronomer, Complejo Astronómico El Leoncito operated under agreement between the Consejo Nacional de Investigaciones Científicas y Técnicas de la República Argentina and the National Universities of La Plata, Córdoba and San Juan.

ization is the principal excitation mechanism. However, these authors suggest that shocks generated by the outflow could provide an additional amount of energy for line formation.

In a study of NLS1 galaxies, a blueshift of $\sim 10 \text{ \AA}$ ($\sim 600 \text{ km s}^{-1}$) was measured in the $[\text{OIII}]\lambda 5007$ emission profile of the galaxy Zw1 (Boroson & Oke 1987). In addition, Aoki et al. (2005) found two NLS1 galaxies with $[\text{OIII}]\lambda 5007$ blueshifts of $\sim 1000 \text{ km s}^{-1}$ relative to $\text{H}\beta$. These objects also present line widths of $1000\text{--}2000 \text{ km s}^{-1}$. These authors suggest that outflows take place in their nuclei and interact with the ambient gas. Bian et al. (2005) studied a sample of 150 Sloan Digital Sky Survey NLS1 galaxies and found 7 blue outlier objects and a correlation between the $[\text{OIII}]\lambda 5007$ blueshift and the FWHM of the core component of $[\text{OIII}]\lambda 5007$ respect to $\text{H}\beta$ in NLS1 galaxies to investigate the dynamics of the narrow-line region (NLR) and to detect possible blue outlier galaxies.

On the other hand, and related to the blue wings, NLS1 galaxies often show blue asymmetric profiles in the $[\text{OIII}]\lambda\lambda 4959, 5007$ lines (e.g., Heckman et al. 1981; Vrtilek & Carleton 1985). These asymmetric profiles are mainly composed by a core component that is emitted from the NLR, which has typical values of FWHM of a few hundred km s^{-1} , and a blue wing with higher FWHM ($\sim 500 \text{ km s}^{-1}$ – 1500 km s^{-1}), which originated closer to the active nucleus (Véron-Cetty et al. 2001; Bian et al. 2005; Komossa et al. 2008). Some studies suggest that the difference of velocities between the core and asymmetric component of $[\text{OIII}]\lambda 5007$ lines correlates with the FWHM of these lines (Cracco et al. 2016; Xiao et al. 2011). In this paper we study the FWHM and luminosity of the blue wing emission and their possible connection with the central black hole.

In order to study the blueshifts and blue wings and their relation with the central engine, we studied 28 NLS1 galaxies with marked blue wings in the $[\text{OIII}]\lambda 5007$ emission line. Of these objects, 22 galaxies lie in the southern hemisphere. We carried out a spectroscopic study of these objects to analyze the outflows and the asymmetries of the profiles. We present the sample and the observations in Sect. 2, the measurement process of the emission lines are described in Sect. 3. In Sect. 4, the blueshifts and blue wings of $[\text{OIII}]\lambda 5007$ and their relation with the black holes are studied. In Sect. 5 we analyze the results and finally, we draw our conclusions in Sect. 6. Throughout this paper, we use the cosmological parameters $H_0 = 70 \text{ Km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$.

2. The sample and observations

For this work, we selected the 20 NLS1 galaxies of the sample of Schmidt et al. (2016), which exhibit marked blue wings in the $[\text{OIII}]\lambda 5007$ emission line. The objects were originally selected from the Véron & Véron catalog (Véron-Cetty & Véron 2010) whose nuclear kinematics have been poorly or even not studied; these objects have redshifts $z < 0.15$ and $\delta \leq 10^\circ$. We added 8 objects from the same catalog, 6 of which have the same characteristics as the main sample and 2 have $z=0.1547$ and $z=0.2063$. The final sample consists of 28 poorly studied NLS1 galaxies (most of them from the southern hemisphere) with an evident asymmetry in the $[\text{OIII}]\lambda 5007$ emission line. Table A.1 lists the galaxy name, right ascension, declination, radial velocity, redshift, apparent magnitude and filter, major diameter, and minor diameter. The data were taken from the Nasa Extragalactic Database (NED). The object MCG–04.24.017 is the only source with morphological classification, as it is classified as a SB0 galaxy (de Vaucouleurs et al. 1991).

Observations were performed in different campaigns between 2011 and 2015 using the REOSC Spectrograph at 2.15 m. telescope of the Complejo Astronómico el LEONCITO (CASLEO), in Argentina. The spectrograph has attached a Tektronix 1024×1024 CCD with $24 \text{ }\mu\text{m}$ pixels. The galaxies were observed using a 2.7 arcsec wide slit and the extractions of each spectrum were of $\sim 2.3 \text{ arcsec}$. We used a 600 line mm^{-1} grating giving a resolution of 190 km s^{-1} FWHM around $\text{H}\beta$. For this paper we obtained spectra in the blue range 4300 \AA to 5200 \AA for the whole sample. In addition, we also obtained spectra for 6 of the additional galaxies in the red range 5800 \AA to 6800 \AA , with a resolution of 170 km s^{-1} FWHM around $\text{H}\alpha$. Table A.2 lists the observational data for the sample, such as the date of the observation and the exposure time of the blue and red spectra. Figure A.1 presents the spectra for the 28 NLS1 galaxies in the blue spectral region 4300 \AA to 5200 \AA . All spectra are in rest frame considering the peak of the $[\text{OIII}]\lambda 5007$ profile. In this paper, we focus on this spectral range in order to study $\text{H}\beta$ and $[\text{OIII}]\lambda 5007$ emission lines in detail. In Fig. A.2, we show the spectra for 6 additional galaxies in the red spectral range 5800 \AA to 6800 \AA . In this way, we can measure emission lines such as $\text{H}\alpha$ and $[\text{NII}]\lambda\lambda 6548, 6584$ and estimate the black hole masses (see Sect. 4.3).

For a mean distance of 240 Mpc (mean redshift of 0.058) the obtained spectra correspond to a central projected distance of 3 kpc. According to some previous studies, the host contribution emission in NLS1 galaxies is around 30% (e.g., Zhou et al. 2006). Therefore, some of the studied emission lines in this work could be contaminated. On the other hand, the outflows seem to be constrained within radial spacial scales of few kpc from the central engine (Villar-Martín et al. 2011). The high excitation suggests that stars are not responsible for ionizing the outflowing gas, but rather AGN related phenomena. These studies claimed that the outflows are within the AGN ionization cones. Rice et al. (2006) suggested that the asymmetries are mainly due to nuclear emitting gas, rather than material distributed throughout the NLR on larger scales. Considering that the broad component of $\text{H}\beta$ is originated in the BLR and taking into account what we mentioned above, the only emission line of interest in our present work that could be contaminated, is the core component of $[\text{OIII}]\lambda 5007$. Following Zhou et al. (2006), this contamination could be around 30%.

3. Measurements

3.1. Iron emission lines

NLS1 galaxies usually exhibit, on average, strong optical Fe emission lines in the vicinity of $\text{H}\beta$ and $[\text{OIII}]\lambda 5007$. These objects exhibit a wider range of iron emission compared with broad line AGNs (e.g., Shen & Ho 2014). According to the lower term of their transitions, these Fe lines are of the S group. This group consists of several lines, which correspond to multiplets 41, 42, 43. The lines of the multiplet 42 overlap with $\text{H}\beta$ and $[\text{OIII}]\lambda\lambda 4959, 5007$ profiles, which are the lines of interest in this work. Because of this, we subtracted the iron emission using the online software¹ developed by Kovačević et al. (2010) and Shapovalova et al. (2012). This software provides a best-fit model that reproduces the iron multiplets in the $\text{H}\beta$ region for each object in function of the gas temperature, Doppler broadening, and shift of the FeII lines.

¹ http://servo.aob.rs/FeII_AGN/

As a first approximation we used the initial parameters provided in the web page. These fits were visually inspected and the initial parameters were modified until acceptable fits were obtained. Some galaxies such as 1RXS J040443.5–295316, 2MASX J01413249–1528016, 2MASXJ21565663–1139314, CTSH34.06, CTSM02.47, and Zw049.106 do not show Fe emission and the fit was not necessary. The Fe multiplet fitting was not satisfactory in RX J0024.7+0820, SDSSJ161227.83+010159.8, and SDSS J225452.22+004631.4, probably due to the low spectral S/N, and we decided not to take this fitting into account. The left panels of Fig. A.1 show the Fe fits on the continuum subtracted spectra of the galaxies.

3.2. Emission line measurements

We are interested in the emission lines, which provide important information about the kinematics of the central regions of AGNs. Emission lines in NLS1 galaxies can be represented by a single or a combination of Gaussian functions (e.g., Xu et al. 2007; Villar-Martín et al. 2011; Schmidt et al. 2016; Scharwächter et al. 2017). For this purpose, we used the LINER routine (Pogge & Owen 1993), which is a χ^2 minimization algorithm that can fit several Gaussian functions to a line profile.

At first, we fitted the continuum in the vicinity of the lines of interest. The fit procedures for the H α + [NII] λ 6548, 6584 profiles, and [SII] λ 6716, 6731 emission lines are detailed in Schmidt et al. (2016).

We fitted the [OIII] λ 5007 emission line with one Gaussian function for the core component ([OIII] $_{cc}$) and with another Gaussian function for the asymmetric component (e.g., Komossa et al. 2008). To fit the narrow component (NC) of H β we had taken into account that it should have approximately the same FWHM as [OIII] $_{cc}$ (e.g., Villar-Martín et al. 2011; Scharwächter et al. 2017). For the broad component (BC) of H β , we fitted one or two Gaussian functions (e.g., Komossa et al. 2008), depending on the case. We followed the same criterion as in Schmidt et al. (2016) of minimizing the residuals obtained with the smaller number of Gaussian functions. All the residuals were visually inspected.

In order to estimate the uncertainties, we measured 15 times some galaxies with different S/N (for example, S/N of around 5, 13, and 22). We assumed that the errors are given by the dispersion in the distribution of the measurements, where we adopted them at 1σ . The relative errors of the flux and FWHM of the broad component of H β and H α are in the ranges 4%–10% and 5%–20%, respectively. The error for the flux and FWHM of the core component of [OIII] are in the ranges 3%–12% and 1%–5%, respectively. For the asymmetric component of [OIII], the relative error is 2%–10% for the flux and 4%–15% for the FWHM. Nevertheless, these uncertainties concern only the measurement process and are a lower limit of the global errors. We did not consider, for example, the uncertainties in the observation system or calibrations. The narrow component of H β and the core component of [OIII] are characterized by high covariance.

As mentioned, strong optical iron emission lines are usually observed in NLS1 spectra, therefore, to see how much the Fe multiplet fitting procedure affects measurements of [OIII] fluxes, FWHM, and asymmetries, we considered some galaxies with an intense iron emission (for example, FAIRALL107, RBS02.19, and RHS56) and we repeated the spectral analysis several times under two different extreme conditions. We fit both the observed and FeII subtracted spectra and we obtained two different mean values for the flux and FWHM of the narrow and the broad component of H β and the core and asymmetric component of [OIII] line. Under this extreme condition and considering high iron

emitting galaxies, we obtained that the variations in the FWHM and flux of the narrow and broad components of H β are less than 5%. The variation of the FWHM and flux of the core component of [OIII] line is 1%–4% and 3%–12%, respectively. As we expected, the higher variation is seen considering the asymmetric component of [OIII] line, being of 5%–10% and 10%–20%, for the FWHM and the flux, respectively.

Figure 1 shows the distributions of FWHM of BC β , NC β , [OIII] $_{cc}$, and [OIII] $_{ac}$. It can be seen that most of the galaxies have FWHM of BC β of ~ 2000 km s $^{-1}$ (top left panel). The top right panel shows that the FWHM of NC β is in the range of 200–800 km s $^{-1}$. The same range of velocities is shown by [OIII] $_{cc}$ in the bottom right panel, and finally, the FWHM of [OIII] $_{ac}$ is in the range of 300–1400 km s $^{-1}$. All values are in agreement with previous measurements of lines emitted at the BLR and NRL (e.g., Osterbrock 1989; Rodríguez-Ardila et al. 2000; Schmidt et al. 2016) and in concordance with the known values of the blue wings of the [OIII] λ 5007 emission lines (e.g., Bian et al. 2005). All the presented FWHM were corrected by the instrumental width as $\text{FWHM}^2 = \text{FWHM}_{obs}^2 - \text{FWHM}_{inst}^2$, where FWHM_{obs} is the measured FWHM and FWHM_{inst} is the instrumental broadening ($\sim 3\text{Å}$ or 190 km s $^{-1}$).

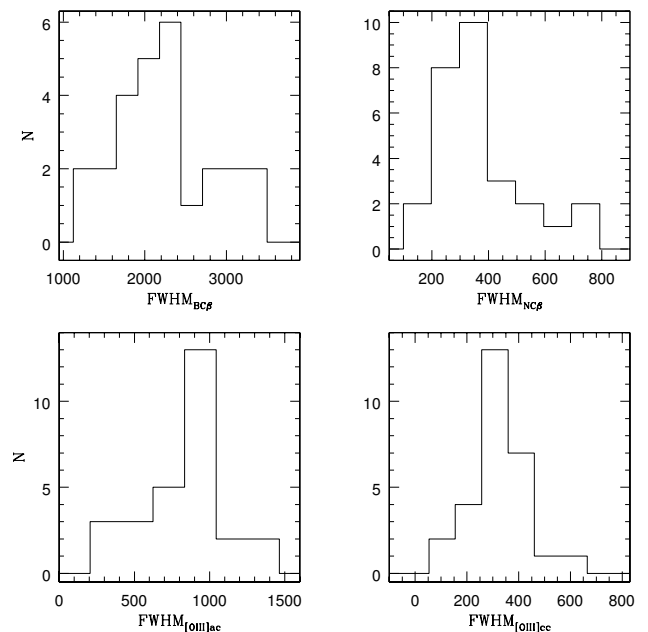


Fig. 1: FWHM distributions of the emission lines in km s $^{-1}$. From left to right and top to bottom: FWHM of BC β , NC β , [OIII] $_{ac}$, and [OIII] $_{cc}$ are shown. All FWHM were corrected by the instrumental width.

4. Emission profile of [OIII] λ 5007

4.1. Blueshifts

In order to determine the blueshift of [OIII], we measured the wavelength difference between the centroid of the narrow component of H β (at $\lambda = 4861.3$ Å) and the core component of [OIII] (at $\lambda = 5006.8$ Å). Hence, the blueshift of [OIII] relative to H β can be calculated based on the laboratory wavelength difference

of both emission lines (145.5 Å). The uncertainties of the determinations of Δv are typically of $\sim 15\%$ – 20% .

In the left panel of Fig. 2, we present the distribution of the [OIII] blueshift, Δv , in units of km s^{-1} , which has a standard deviation of 137 km s^{-1} and an interquartile range (IQR) of 104 km s^{-1} . According to the criteria adopted by Zamanov et al. (2002), there is one blue outlier galaxy with $\Delta v < -250 \text{ km s}^{-1}$. This object is V961349–439 and has a value of Δv of $-405 \pm 47 \text{ km s}^{-1}$. This object presents two blueshifted broad components of $\sim 3100 \text{ km s}^{-1}$ and 1300 km s^{-1} and one narrow component of 360 km s^{-1} in the $\text{H}\beta$ emission profile. On the other hand, the [OIII] $\lambda\lambda 4959, 5007$ lines of this galaxy show a core component of $\sim 360 \text{ km s}^{-1}$ and a blueshifted component of $\sim 680 \text{ km s}^{-1}$.

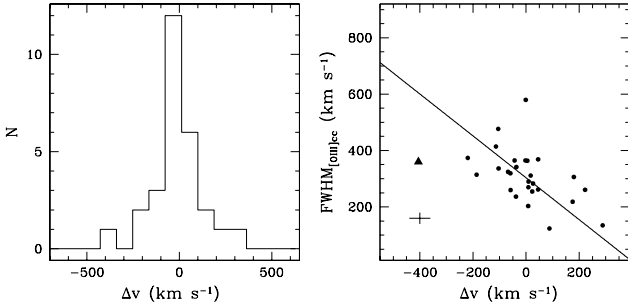


Fig. 2: Left: Histogram of the [OIII] blueshift relative to $\text{H}\beta$ (Δv) in units of km s^{-1} . Right: The relation between the [OIII] blueshift and the FWHM of the core component of [OIII] in units of km s^{-1} is shown. The blue outlier galaxy V961349–439 is indicated by the triangle. The typical error bar is shown. The solid line represents the best fit for our data.

Galaxies with larger [OIII] blueshifts tend to present broader FWHM for the core component of [OIII]. That can be seen in the right panel of Fig. 2, which shows the relation between Δv and the FWHM of [OIII] $_{cc}$, which have a modest Pearson coefficient correlation of $r_p = -0.48$ and a p-value of 9×10^{-3} . This agrees with previous results of Bian et al. (2005) and Komossa et al. (2008). The blue outlier galaxy V961349–439 is indicated in the following plots as a triangle. The solid line in the plot indicates the OLS bisector fit for our data, which is $\text{FWHM}_{[\text{OIII}]_{cc}} = (303 \pm 18) \text{ km s}^{-1} - (0.74 \pm 0.16) \Delta v (\text{km s}^{-1})$.

4.2. Blue wings

As was previously mentioned, all the galaxies of the sample present asymmetries (most of which are very evident) in the [OIII] $\lambda 5007$ emission profiles and were fitted with an additional Gaussian component (Sect. 3.2). As an example, Fig. 3 shows the fits of the [OIII] $\lambda 5007$ line for two galaxies, where the composite profile is shown with a thick solid line, the core and asymmetric components are shown with a dotted line, and the residuals of the fit are plotted below.

We defined the parameter Δv_{asym} as the difference of velocities of the centroid of [OIII] $_{ac}$ and that of [OIII] $_{cc}$. Typical uncertainties in the measurements of Δv_{asym} are on the order of $\sim 15\%$ – 20% .

The left panel of Fig. 4 shows the distribution of Δv_{asym} , which has a standard deviation of 207 km s^{-1} and IQR of 187 km s^{-1} . The range of the asymmetric emission of [OIII] $\lambda 5007$ is in

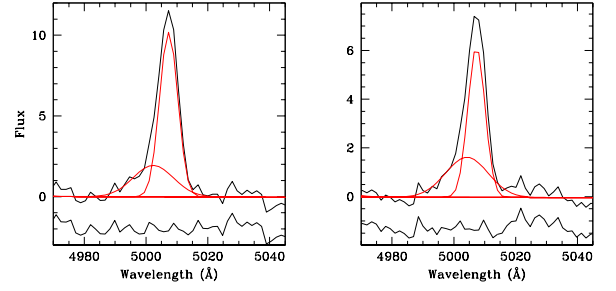


Fig. 3: Example of the two Gaussian fit for the [OIII] $\lambda 5007$ profile of the galaxies Zw049.106 (left) and SDSSJ134524.70–025939.8 (right). Thick line represents the spectra; the core and asymmetric components of the [OIII] line are shown with dotted lines. The fit (sum of the two components) is superimposed on the spectra. The residuals are plotted at the bottom of each panel. The flux is in arbitrary units in both plots.

agreement with previous known results (e.g., Véron-Cetty et al. 2001). There are some galaxies with very marked asymmetries ($\Delta v_{\text{asym}} < -250 \text{ km s}^{-1}$): IRAS20520–2329 ($\Delta v_{\text{asym}} = -267 \text{ km s}^{-1}$), Zw049.106 ($\Delta v_{\text{asym}} = -302 \text{ km s}^{-1}$), CTSM02.47 ($\Delta v_{\text{asym}} = -308 \text{ km s}^{-1}$), RXJ0902.5–0700 ($\Delta v_{\text{asym}} = -331 \text{ km s}^{-1}$), 2MASXJ01413249–1528016 ($\Delta v_{\text{asym}} = -367 \text{ km s}^{-1}$), WPV85007 ($\Delta v_{\text{asym}} = -432 \text{ km s}^{-1}$), IRAS04576+0912 ($\Delta v_{\text{asym}} = -498 \text{ km s}^{-1}$), RBS0219 ($\Delta v_{\text{asym}} = -566 \text{ km s}^{-1}$), RXJ0024.7+0820 ($\Delta v_{\text{asym}} = -711 \text{ km s}^{-1}$), and RXJ2301.8–5508 with $\Delta v_{\text{asym}} = -893 \text{ km s}^{-1}$ (see Table A.4).

The galaxies with lower values of Δv_{asym} , also tend to present higher FWHM of [OIII] $_{cc}$, in agreement with recent results (Cracco et al. 2016). This is shown in the right panel of Fig. 4. This relation also has a modest Pearson correlation coefficient of $r_p = -0.59$ with a p-value $< 10^{-3}$. The solid line in the plot represents the OLS bisector fit for our sample, which is $\text{FWHM}_{[\text{OIII}]_{cc}} = (169 \pm 29) \text{ km s}^{-1} - (0.63 \pm 0.14) \Delta v_{\text{asym}} (\text{km s}^{-1})$.

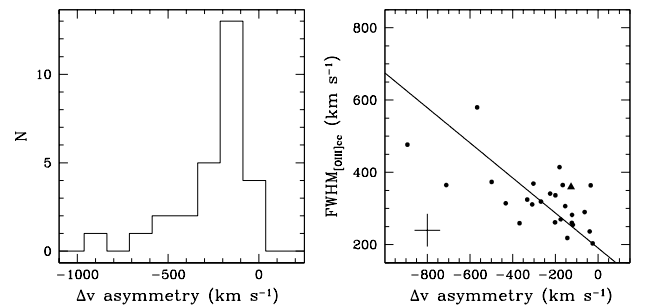


Fig. 4: Histogram of Δv_{asym} (left). Relation between Δv_{asym} and $\text{FWHM}_{[\text{OIII}]_{cc}}$ (right) is shown. Both parameters are in units of km s^{-1} . The triangle indicates the blue outlier galaxy, V961349–439. The solid line represents the best fits for our data. A typical error bar is shown in the plot.

No correlation is seen considering Δv_{asym} and the FWHM of the remaining emission lines. On the other hand, we can not say that there is a correlation between Δv_{asym} and the [OIII] blueshift Δv either. The two parameters seem to be weakly correlated, given we find a Pearson correlation coefficient of $r_p = 0.33$ and a p-value of 0.08, possibly because of the small number of objects.

4.3. Total asymmetry

To study the total asymmetric emission, that is, the total extension of the blue wing, in addition to Δv_{asym} we have to consider half of the width of the asymmetric component in the base, i.e.,

$$\text{total asymmetry} = \Delta v_{\text{asym}} - \frac{1}{2} \text{Width}_{\text{base}}[\text{OIII}]_{\text{ac}}. \quad (1)$$

Although it is difficult to estimate the width of the base, we can assume that $\text{Width}_{\text{base}}[\text{OIII}]_{\text{ac}}$ is $\sim 2 \text{ FWHM}[\text{OIII}]_{\text{ac}}$, then we have

$$\text{total asymmetry} = \Delta v_{\text{asym}} - \text{FWHM}[\text{OIII}]_{\text{ac}}. \quad (2)$$

In this way, the thus defined total asymmetry parameter is a good estimator of the radial velocity of the extreme of the blue wing relative to the centroid of the core component of [OIII] λ 5007. As mentioned in Sect 3.2, to estimate the uncertainties of the measurements, we repeated the measurement procedure several times in galaxies with different S/N. In this manner, the relative error of the total asymmetry is typically 4%–15%.

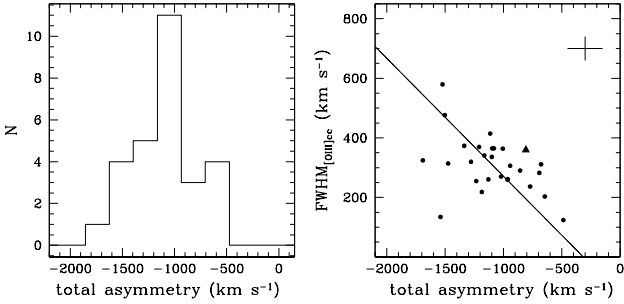


Fig. 5: Histogram of the total asymmetry, given by equation 2 (left). The relation between the total asymmetry and $\text{FWHM}_{[\text{OIII}]_{\text{ac}}}$ (right) is shown. Both parameters are in units of km s^{-1} . The triangle indicates the blue outlier galaxy, V961349–439. The solid line represents the best fit for our data. A typical error bar is shown.

The left panel of Fig. 5 shows the distribution of the total asymmetry, in units of km s^{-1} . The galaxies of our sample have [OIII] blue wings with an extension (relative to the centroid of [OIII] $_{\text{cc}}$) that ranges from $\sim -1700 \text{ km s}^{-1}$ to $\sim -500 \text{ km s}^{-1}$; most galaxies show blue wings of $\sim -1100 \text{ km s}^{-1}$. The distribution has a standard deviation of 298 km s^{-1} and IQR of 375 km s^{-1} . Just like Δv and Δv_{asym} , the total blue wing also correlates with the FWHM of [OIII] $_{\text{cc}}$ (right panel). This tendency has a modest Pearson correlation coefficient of $r_p = -0.46$ and a p-value of 0.01. The solid line represents the OLS bisector fit for our data, which is $\text{FWHM}_{[\text{OIII}]_{\text{ac}}} = (-125 \pm 113) \text{ km s}^{-1} - (0.40 \pm 0.09) \times \text{total asymmetry} (\text{km s}^{-1})$.

On the other hand, to study the relation between the blue wing and central engine, it is necessary to estimate the black hole mass. For the 28 galaxies of our sample, 20 of these sources already have a black hole mass estimation in Schmidt et al. (2016). We calculated the black hole mass of the 8 remaining galaxies, using equation 6 of Greene & Ho (2005). For 6 of the galaxies, we used the FWHM and luminosity of the broad component of H α because the black hole mass estimation through this line is more reliable than the estimation through H β because of its higher S/N ratio (Greene & Ho 2005). For the 2 remaining galaxies, the parameters of the broad component of H β were used through equation 7 of the mentioned work.

The used formalism in the black hole estimation assumes the correlation between the size of the BLR and the luminosity at 5100 \AA (see Greene & Ho 2005). According to some authors, for highly accreting sources, this scaling relation seems to deviate from those of lower accretion rate (e.g., Wang et al. 2014a,b). Related to this, for a given L_{5100} , a high accretion source has shorter H β lags by a factor of up to $\sim 3 - 4$ compared with lower accretion rate objects (Du et al. 2015), indicating a smaller size of the BLR. This difference in the estimation of the BLR size for high accreting galaxies produces a black hole mass uncertainty of 0.5 dex.

Table A.3 presents the FWHM and luminosity of BC α and the black hole masses for 8 galaxies. They are in the same range of black hole mass for NLS1 galaxies previously studied (e.g., Grupe 2004; Schmidt et al. 2016; Järvelä et al. 2017).

Analyzing the total asymmetry, we found a very interesting correlation between this asymmetry and the black hole mass; we find a Pearson coefficient of $r_p = -0.63$ and a p-value of 3×10^{-4} . This tendency is stronger than the correlation between M_{BH} and $\text{FWHM}_{[\text{OIII}]_{\text{ac}}}$ ($r_p = 0.53$) and that between M_{BH} and Δv_{asym} ($r_p = -0.20$). Considering galaxies with blue asymmetries in the [OIII] λ 5007 profile, objects with higher black hole masses tend to show more extended blue wings (left panel of Fig. 6).

On the other hand, the black hole mass also correlates with the luminosity of [OIII] $_{\text{ac}}$ (middle panel). This tendency has a moderate Pearson correlation coefficient of $r_p = -0.68$ and a p-value of 7×10^{-5} . Galaxies with larger black hole masses seem to present higher luminosities of the blue wing. The solid line in the plot represents the OLS bisector fit for our data, which is $\log(L_{[\text{OIII}]_{\text{ac}}}) = (33 \pm 1) + (1.13 \pm 0.15) \times \log(M_{\text{BH}}/M_{\odot})$.

If the total asymmetry correlates with the central black hole mass, we can also expect a correlation between the former and the luminosity of the broad component of H β . This tendency can be seen in the right panel of Fig. 6, where the solid line represents the best OLS bisector fit for the data. This relation has a Pearson coefficient of $r_p = -0.43$ and a p-value < 0.03 .

The Eddington luminosity, given by $L_{\text{Edd}} = 1.26 \times 10^{38} M_{\text{BH}}/M_{\odot} \text{ ergs}^{-1}$, is closely linked to black hole mass. We calculated the Eddington ratio between the bolometric luminosity (L_{bol}) and the Eddington luminosity, considering $L_{\text{bol}} \sim 3500 \times L_{[\text{OIII}]}$ (Heckman et al. 2004). Assuming that the uncertainty of the Eddington ratio is given by the error of the luminosity of [OIII] and the uncertainty of the black hole mass estimation, the resulting lower limit for the error of the Eddington ratio is ~ 0.5 dex.

As expected, the luminosities of some objects in our sample are beyond the Eddington limit, which is in agreement with previous results (e.g., Cracco et al. 2016; Xu et al. 2007; Bian & Zhao 2004). Considering the Eddington ratio and Δv , no significant correlation was found between both parameters, with a Pearson

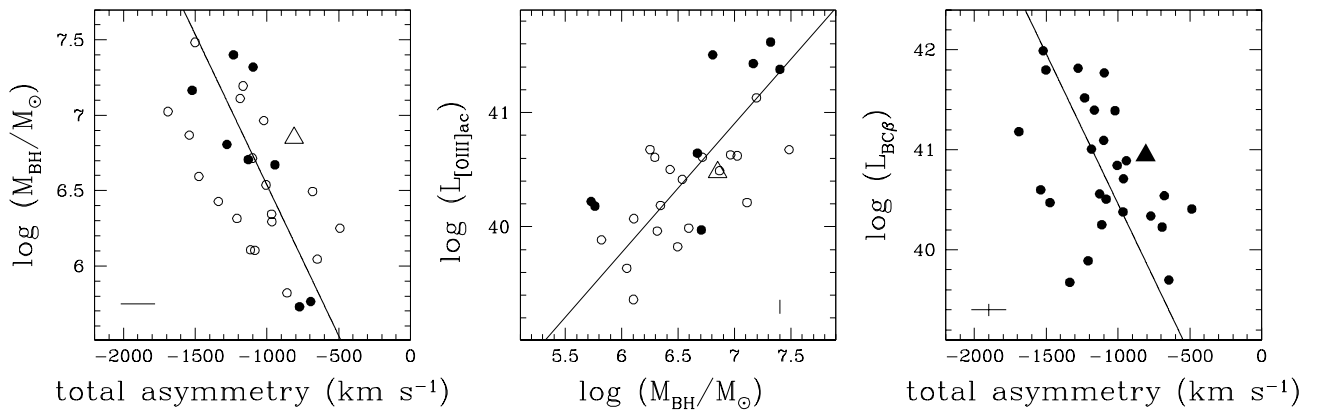


Fig. 6: Left: Relation between the total asymmetry (given by equation 2), in units of km s⁻¹, and the black hole mass (in units of M_⊙). Middle: The correlation among the black hole mass and the luminosity of the blue wing, [OIII]_{ac} (in units of erg s⁻¹) is shown. Right: The tendency between the total asymmetry (in km s⁻¹) and the luminosity of the broad component of Hβ (in units of erg s⁻¹) is shown. Open symbols represent galaxies whose black hole mass estimations were extracted from Schmidt et al. (2016). The triangle indicates the blue outlier galaxy V961349–439. The solid lines represent the best fit for our data. Typical error bars are shown.

correlation coefficient of $r_p = -0.13$ and a p-value of 0.5; this agrees with results of Cracco et al. (2016) and Bian et al. (2005). In addition, given a Pearson coefficient of $r_p = -0.09$, no correlation was found between the Eddington ratio and Δv_{asym} . On the other hand, a weak correlation was found between the Eddington ratio and the total asymmetry, with the Pearson coefficient of $r_p = 0.33$ and a p-value of 0.08. Also, an anti-correlation between the Eddington ratio and the black hole mass is observed and the correlation coefficient of $r_p = -0.42$ and p-value < 0.03.

Table A.4 presents the FWHM and the flux of the core and the asymmetric emission of [OIII]λ5007, the values of the [OIII] blueshift (Δv), the amount of asymmetry (Δv_{asym}), the total asymmetry (equation 2) and the Eddington ratio. All presented FWHM were corrected by the instrumental broadening, as was mentioned in Sect. 3.2.

As we said in Sect. 2, 20 of the 28 studied objects in this work were taken from the sample of 36 galaxies of Schmidt et al. (2016) because they present asymmetric [OIII] profiles. The black hole mass range of galaxies that show asymmetric profiles is almost the same as that of galaxies that show symmetric lines. This means that blue wings can be present in different galaxies, independently of the black hole mass. In this scenario, galaxies with a large black hole mass may show non-asymmetric [OIII] emission profiles and galaxies with lower black hole masses may exhibit asymmetric lines. Related to this, 94% of the sample of NLS1 of Cracco et al. (2016) show asymmetric [OIII] emission profiles, and these authors have claimed that no difference appears to exist between NLS1 and BLS1, despite the difference in the black hole masses. In order to try to detect any possible difference between galaxies that show asymmetric profiles and objects that do not, we checked the Eddington ratio. Both groups of galaxies present almost the same range of Eddington ratio.

Although we do not know which mechanisms trigger the outflows, when we consider only galaxies with asymmetric [OIII] emission profiles, we found that there is a correlation between the total asymmetry and the black hole mass. Nevertheless, there is no correlation between these parameters when considering a sample that includes symmetric emitting galaxies.

5. Discussion

In a study of the presence of winds in a large spectroscopic sample of around 600000 local galaxies, Concas et al. (2017) found that in star forming galaxies, the ionized interstellar gas traced by the [OIII]λ5007 line never appears to be outflowing. In these galaxies, the [OIII] line profile is perfectly fitted by a single Gaussian without need of a second component. These authors claimed that the need for a second broader Gaussian component increases with a clear trend with the increase of the AGN contribution to the galaxy spectrum. In this scenario, blue wings were found in 73% of a sample of intermediate-type Seyfert and in 68% of a sample of Seyfert 2 galaxies (Vaona et al. 2012). Woo et al. (2016) studied the outflows in a sample of type 2 AGNs and they found that 43% of the galaxies present winds. These decreasing percentages, going from type 1 to type 2 AGNs, are in agreement with the results of Concas et al. (2017). This is in compliance with the fact that asymmetric profiles are usually detected in quasars, Seyfert galaxies, and in particular are very common in NLS1 (e.g., Véron-Cetty et al. 2001; Xiao et al. 2011). In a sample of 296 NLS1, Cracco et al. (2016) found that 94% of the galaxies exhibit [OIII] emission profiles with asymmetries.

In our sample of 28 asymmetric-line emitting NLS1 galaxies, the two components of the [OIII]λ5007 profile show different line widths and have velocities that range between 200 and 800 km s⁻¹ for the core component and between 300 and 1400 km s⁻¹ for the asymmetric component. This can be interpreted as two kinematically distinct regions, which in agreement with some authors (e.g., Holt et al. 2003; Bian et al. 2005). The range of velocities of the asymmetric component of [OIII] has values between the one of the core component and the range of the BLR (Sect. 3.2). If we assume that the square of the FWHM of the gas decreases with the distance to the center (e.g., Schmidt et al. 2016), one could expect that the asymmetric component is originated at the inner parts of the NLR between the emitting region of the core component of [OIII]λ5007 line and the BLR (by definition, the emitting region of the broad component of Hβ).

This indicates that the kinematic of the gas in the inner parts of the NLR is more turbulent compared to the external emitting regions, in agreement with several studies (e.g., Véron-Cetty et al. 2001; Bian et al. 2005). Related to this, Cracco et al. (2016) suggested that the asymmetry is probably caused by the presence of outflowing gas from the inner regions of the active nucleus, which interacts with the surrounding medium by transferring kinetic energy and by reducing the equivalent width of the [OIII] line as its velocity increases. This is in agreement with the results of Ludwig et al. (2012), who found that the shift of the blue wing correlates with the equivalent width of the [OIII] emission.

The blue outlier galaxy V961349–439 does not deviate in any of the studied relations. These tendencies mainly involve parameters of the blue wing such as the total asymmetry, the luminosity of the asymmetric component, and Δv_{asym} . There is no significant correlation between the mentioned parameters and the blueshift Δv as the Pearson coefficient values are between ~ 0.1 and ~ 0.3 . This is in agreement with the idea that the [OIII] blueshift is not connected to the blue wing (Boroson 2005; Cracco et al. 2016). This could suggest that the mechanism that produces the blueshift is not the same as that yielding the asymmetry of the profile.

Considering the case of V961349–439, this object presents a higher [OIII] blueshift relative to $H\beta$ ($\Delta v = -405 \text{ km s}^{-1}$), but shows a blue wing with an extension of $\sim -810 \text{ km s}^{-1}$, which is one of the lower values of the sample; this is in concordance with the analysis above.

However, it has been proposed that the blueshift of [OIII] is the result of an outflowing gas from the central regions (Zamanov et al. 2002). The outflowing gas could be originated at the inner NLR and is possibly related to the wind of the nucleus (Elvis 2000). In this scenario, both the blueshift and the blue wings would be originated at the same region.

Aside from the [OIII] emission, there are high ionization emission lines emitted at the NLR, which also show shifts, as is the case of [Fe VII] (Rodríguez-Ardila et al. 2006; Cracco et al. 2016), [Ne III] and [Ne V] (Spoon & Holt 2009), and [Fe VII], [Fe X] and [Fe XI] (Rodríguez-Ardila et al. 2006; Komossa et al. 2008). Moreover, the high ionization C IV λ 1549 line, emitted from the BLR, was also found to be blueshifted (e.g., Sulentic et al. 2007; Kovačević et al. 2007). In this scenario, and considering that the higher the ionization potential, the line originates in more internal regions, the outflow would be moving through a stratified medium ionized by the central engine (e.g., Cracco et al. 2016).

On the other hand, the found correlation between the black hole mass and the total asymmetry, and the relation between the former and the luminosity of the asymmetric component of [OIII] line, suggest that the winds originated at the NLR and manifested through the [OIII] $_{\text{ac}}$ emission are affected by the central engine. This would be in agreement with the fact that the black hole mass correlates with the radio emission, suggesting that NLS1 with larger black hole masses are more likely to show relativistic jets (Järvelä et al. 2015) and therefore, strong winds. In addition, the correlation between the total asymmetry and the black hole mass is in agreement with the idea that wider extended components could be explained by virial effects of the central black holes, as suggested by Zhang & Feng (2017).

We found that the Eddington ratios of the galaxies of the sample are in the same range as the obtained by other authors (Xu et al. 2007; Bian & Zhao 2004); some objects show luminosities close to the Eddington luminosity, as it is expected in NLS1 galaxies (e.g., Cracco et al. 2016; Xu et al. 2007).

There is a mild correlation between the total asymmetry and the Eddington ratio, for which we find a Pearson correlation coefficient of $r_p = -0.33$. We also found trends between the total asymmetry and the luminosities of the broad component of $H\beta$ and [OIII] λ 5007 with $r_p = -0.43$ and $r_p = -0.40$, respectively. The fact that the total asymmetry correlates with the Eddington ratio, the emission of the BLR, and the emission of [OIII] could indicate that the extension of the blue wing also depends on the intensity of the AGN; this is in agreement with previous results of Mullaney et al. (2013) and Woo et al. (2016) in type 2 AGNs. Nevertheless, these correlations are not so marked as the relation between the total asymmetry and the black hole mass, which has a Pearson correlation coefficient of $r_p = -0.63$. This implies, as is expected, that asymmetric profiles are more related to kinematical processes than to photoionization mechanisms.

Related to the [OIII] profile, some studies found that the bulge of the host galaxy has influence in the width of the emission lines emitted at the NLR, i.e., the [OIII] λ 5007 line (e.g., Nelson & Whittle 1996). As we showed, the asymmetric emission of [OIII] has a marked relation with the central engine, not only in the shape, but also in the luminosity. Thus, it could be said that the [OIII] λ 5007 emission profile is doubly affected, by the bulge of the host galaxy and by the central engine of the AGN.

6. Final remarks

We analyzed a sample of 28 poorly studied NLS1 galaxies, most of which are from the southern hemisphere. All of these galaxies present a blue asymmetric component in the [OIII] λ 5007 emission line. After a careful Gaussian decomposition, we measured the radial velocity difference between the core component of [OIII] λ 5007 and the narrow component of $H\beta$. We found that many of the galaxies present a blueshift of [OIII] relative to $H\beta$. The distribution of Δv in our sample has a standard deviation of 137 km s^{-1} and an IQR of 104 km s^{-1} . Considering the criterion adopted by Zamanov et al. (2002), we found one blue outlier galaxy with a blueshift $< -250 \text{ km s}^{-1}$. This object is V961349–439 and has a blueshift of $-405 \pm 47 \text{ km s}^{-1}$. We confirmed a previously known correlation between the [OIII] blueshift relative to $H\beta$ and the FWHM of the core component of [OIII] λ 5007 (Bian et al. 2005; Komossa et al. 2008), which for our sample has a Pearson correlation coefficient of $r_p = -0.48$ and a p-value of 9×10^{-3} . Galaxies with larger blueshift tend to present broader FWHM in the core component of [OIII].

Aside from the [OIII] blueshift relative to $H\beta$, we also analyzed the blue wings of the asymmetric emission of the [OIII] λ 5007 emission profile. We studied the radial velocity difference between the centroid of the asymmetric component relative to the core component, that is Δv_{asym} . We found that the distribution of Δv_{asym} has a standard deviation of 207 km s^{-1} , an IQR of 187 km s^{-1} , and a mean value of -248 km s^{-1} . The range of Δv_{asym} found in this work is in agreement with previous results (e.g., Véron-Cetty et al. 2001; Cracco et al. 2016). The FWHM of the core component of [OIII] λ 5007 also correlates with Δv_{asym} , given a Pearson correlation coefficient of $r_p = -0.59$ and a p-value $< 10^{-3}$. This is in agreement with recent results of Cracco et al. (2016), who found that these parameters are related with a Spearman coefficient of $r_s = -0.46$ and a p-value of 1.4×10^{-18} . Considering Δv_{asym} and the [OIII] blueshift Δv , a very weak tendency seems to relate both parameters; we find a Pearson correlation coefficient of $r_p = 0.33$ and

a p-value of 0.08, which is probably due to the small number of galaxies.

Finally, we studied the radial velocity difference between the blue end of the blue wing and the centroid of the core component of [OIII] λ 5007. In our sample, this total asymmetry (defined by equation 2) ranges from ~ -1700 km s $^{-1}$ to ~ -500 km s $^{-1}$, with most galaxies showing blue wings which extend up to ~ -1100 km s $^{-1}$. This distribution has a standard deviation of 298 km s $^{-1}$ and an IQR of 375 km s $^{-1}$. The total blue wing also correlates with the FWHM of [OIII] $_{cc}$, just like Δv and Δv_{asym} . This tendency has a Pearson correlation coefficient of $r_p = -0.46$ and a p-value of 0.01. Given a Pearson correlation coefficient of $r_p = -0.63$ and a p-value of 3×10^{-4} , we found that the radial velocity of the blue end of the asymmetry relative to the core component of [OIII] correlates with the black hole mass. Galaxies with higher black hole masses tend to show a more marked asymmetry in the [OIII] λ 5007 emission profile, which agrees with the idea of that NLS1 galaxies with larger black hole masses have higher probabilities of showing relativistic jets, and therefore, strong winds (Järvelä et al. 2015).

We calculated the Eddington ratio for the whole sample. As expected, the luminosities of some galaxies are beyond the Eddington limit, in agreement with previous results (e.g., Cracco et al. 2016; Xu et al. 2007; Bian & Zhao 2004). We see a weak correlation between the Eddington ratio and the total asymmetry, with a Pearson correlation coefficient of $r_p = 0.33$ and a p-value of 0.08.

We found a correlation between the luminosity of the asymmetric component of [OIII] and the black hole mass, which has a moderate Pearson correlation coefficient of $r_p = -0.68$ and a p-value of 7×10^{-5} . Related to this, galaxies with larger black hole masses seem to show higher luminosities of the blue wing. Finally, we found a correlation between the luminosity of the broad component of H β and the total asymmetry, with a Pearson coefficient of $r_p = -0.43$ and a p-value $< 3 \times 10^{-2}$. This suggests that the blue wing is influenced by the intensity of the AGN, which is in agreement with previous results (Mullaney et al. 2013; Woo et al. 2016).

Acknowledgements. E.O.S and G.A.O want to thank the considerable help of Damian Mast and Jose Ovejero. We also appreciate the helpful comments and suggestions made by the anonymous referee, which improved this article. This work was partially supported by Consejo de Investigaciones Científicas y Técnicas (CONICET) and Secretaría de Ciencia y Técnica de la Universidad Nacional de Córdoba (SecyT). This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

References

Aoki, K., Kawaguchi, T., & Ohta, K. 2005, *ApJ*, 618, 601
 Bian, W., Yuan, Q., & Zhao, Y. 2005, *MNRAS*, 364, 187
 Bian, W. & Zhao, Y. 2004, *MNRAS*, 352, 823
 Boroson, T. 2005, *AJ*, 130, 381
 Boroson, T. & Oke, J. 1987, *PASP*, 99, 809
 Chelouche, D. & Netzer, H. 2005, *ApJ*, 625, 95
 Churazov, E., Brügggen, M., Kaiser, C., Böhringer, H., & Forman, W. 2001, *ApJ*, 554, 261
 Colbert, E., Baum, S., Gallimore, J., et al. 1996, *ApJS*, 105, 75
 Concas, A., Popesso, P., Brusa, M., et al. 2017, *A&A*, 606, A36
 Cracco, V., Ciroi, S., Berton, M., et al. 2016, *MNRAS*, 462, 1256
 Crenshaw, D. & Kraemer, S. 2007, *ApJ*, 659, 250
 Das, V., Crenshaw, D., Hutchings, J., et al. 2005, *AJ*, 130, 945

de Vaucouleurs, G., de Vaucouleurs, A., Corwin, Jr., H., et al. 1991, *Third Reference Catalogue of Bright Galaxies. Volume I: Explanations and references. Volume II: Data for galaxies between 0 h and 12 h . Volume III: Data for galaxies between 12 h and 24 h .*
 Du, P., Hu, C., Lu, K.-X., et al. 2015, *ApJ*, 806, 22
 Elvis, M. 2000, *ApJ*, 545, 63
 Gallimore, J., Axon, D., O’Dea, C., Baum, S., & Pedlar, A. 2006, *AJ*, 132, 546
 Goodrich, R. 1989, *ApJ*, 342, 224
 Greene, J. & Ho, L. 2005, *ApJ*, 630, 122
 Grupe, D. 2004, *AJ*, 127, 1799
 Grupe, D. & Leighly, K. 2002, in *X-ray Spectroscopy of AGN with Chandra and XMM-Newton*, ed. T. Boller, S. Komossa, S. Kahn, H. Kunieda, & L. Gallo, 287
 Grupe, D., Thomas, H.-C., & Leighly, K. 2001, *A&A*, 369, 450
 Heckman, T., Miley, G., van Breugel, W., & Butcher, H. 1981, *ApJ*, 247, 403
 Heckman, T. M., Kauffmann, G., Brinchmann, J., et al. 2004, *ApJ*, 613, 109
 Holt, J., Tadhunter, C. N., & Morganti, R. 2003, *MNRAS*, 342, 227
 Järvelä, E., Lähteenmäki, A., & León-Tavares, J. 2015, *A&A*, 573, A76
 Järvelä, E., Lähteenmäki, A., Lietzen, H., et al. 2017, *A&A*, 606, A9
 Komossa, S., Xu, D., Zhou, H., Storchi-Bergmann, T., & Binette, L. 2008, *ApJ*, 680, 926
 Kovačević, J., Bachev, R., Popović, L. Č., Zamanov, R., & Marziani, P. 2007, in *American Institute of Physics Conference Series*, Vol. 938, *Spectral Line Shapes in Astrophysics*, ed. L. C. Popovic & M. S. Dimitrijevic, 104–108
 Kovačević, J., Popović, L., & Dimitrijević, M. 2010, *ApJS*, 189, 15
 Krongold, Y., Nicastro, F., Elvis, M., et al. 2007, *ApJ*, 659, 1022
 Ludwig, R. R., Greene, J. E., Barth, A. J., & Ho, L. C. 2012, *ApJ*, 756, 51
 Marziani, P., Sulentic, J., Zamanov, R., et al. 2003a, *ApJS*, 145, 199
 Marziani, P., Zamanov, R., Sulentic, J., & Calvani, M. 2003b, *MNRAS*, 345, 1133
 Moll, R., Schindler, S., Domainko, W., et al. 2007, *A&A*, 463, 513
 Mullaney, J. R., Alexander, D. M., Fine, S., et al. 2013, *MNRAS*, 433, 622
 Nelson, C. & Whittle, M. 1996, *ApJ*, 465, 96
 Osterbrock, D. 1989, *Astrophysics of gaseous nebulae and active galactic nuclei*
 Osterbrock, D. & Pogge, R. 1985, *ApJ*, 297, 166
 Pogge, R. & Owen, J. 1993, *OSU Internal Report 93-01*
 Rice, M. S., Martini, P., Greene, J. E., et al. 2006, *ApJ*, 636, 654
 Rodríguez-Ardila, A., Binette, L., Pastoriza, M., & Donzelli, C. 2000, *ApJ*, 538, 581
 Rodríguez-Ardila, A., Prieto, M., Viegas, S., & Gruenwald, R. 2006, *ApJ*, 653, 1098
 Scharwächter, J., Husemann, B., Busch, G., Komossa, S., & Dopita, M. A. 2017, *ApJ*, 848, 35
 Schmidt, E., Ferreira, D., Vega Neme, L., & Oio, G. 2016, *A&A*, 596, A95
 Shapovalova, A., Popović, L., Burenkov, A., et al. 2012, *ApJS*, 202, 10
 Shen, Y. & Ho, L. C. 2014, *Nature*, 513, 210
 Spoon, H. W. W. & Holt, J. 2009, *ApJ*, 702, L42
 Sulentic, J., Bachev, R., Marziani, P., Negrete, C., & Dultzin, D. 2007, *ApJ*, 666, 757
 Vaona, L., Ciroi, S., Di Mille, F., et al. 2012, *MNRAS*, 427, 1266
 Véron-Cetty, M.-P. & Véron, P. 2010, *A&A*, 518, A10
 Véron-Cetty, M.-P., Véron, P., & Gonçalves, A. 2001, *A&A*, 372, 730
 Villar-Martín, M., Humphrey, A., Delgado, R. G., Colina, L., & Arribas, S. 2011, *MNRAS*, 418, 2032
 Vrtillek, J. & Carleton, N. 1985, *ApJ*, 294, 106
 Wang, J.-M., Du, P., Hu, C., et al. 2014a, *ApJ*, 793, 108
 Wang, J.-M., Du, P., Li, Y.-R., et al. 2014b, *ApJ*, 792, L13
 Williams, R., Pogge, R., & Mathur, S. 2002, *AJ*, 124, 3042
 Woo, J., Bae, H., Son, D., & Karouzos, M. 2016, *ApJ*, 817, 108
 Xiao, T., Barth, A., Greene, J., et al. 2011, *ApJ*, 739, 28
 Xu, D., Komossa, S., Zhou, H., Wang, T., & Wei, J. 2007, *ApJ*, 670, 60
 Zakamska, N., Hamann, F., Pâris, I., et al. 2016, *MNRAS*, 459, 3144
 Zamanov, R., Marziani, P., Sulentic, J., et al. 2002, *ApJ*, 576, L9
 Zhang, X. & Feng, L. 2017, *MNRAS*, 468, 620
 Zhou, H., Wang, T., Yuan, W., et al. 2006, *ApJS*, 166, 128

Appendix A: Additional data

In this Appendix, we present the principal characteristic of the galaxies of the sample (Table A.1), the observational data (Table A.2), black hole masses for eight galaxies (Table A.3), and the parameters of [OIII] λ 5007 (Table A.4). We also show the blue spectra of the galaxies in Fig. A.1 and the red spectra of six galaxies in Fig. A.2.

Table A.1: Sample. Column 1: Galaxy name. Columns 2 to 8: Right ascension (J2000), declination (J2000), radial velocity, redshift, apparent magnitude and filter, major diameter, minor diameter, as taken from NED. Galaxies labeled with "*" were taken from the sample of Schmidt et al. (2016).

Galaxy name	RA J(2000.0)	Dec J(2000.0)	Radial vel. (km s ⁻¹)	Redshift	Mag & Filter	Major Diam (arc min)	Minor Diam (arc min)
1RXSJ040443.5–295316*	04 04 43.28	–29 53 23.1	17995	0.060025	17.3V	0.20	0.19
2MASXJ01413249–1528016*	01 41 32.50	–15 28 01.5	24304	0.081069	16.9V	0.34	0.24
2MASXJ05014863–2253232*	05 01 48.63	–22 53 23.2	12232	0.040800	14.0V	0.31	0.29
2MASXJ08173955–0733089	08 17 39.58	–07 33 08.8	21765	0.072600	16.9R	—	—
2MASXJ21531910–1514111*	21 53 19.13	–15 14 12.1	23324	0.077800	14.7V	—	—
2MASXJ21565663–1139314*	21 56 56.61	–11 39 30.6	8420	0.028086	15.4V	—	—
CTSH34.06	06 09 17.48	–56 06 58.4	9535	0.031805	16.2V	0.45	—
CTSM02.47*	10 46 23.55	–30 04 21.0	17106	0.057060	17.0V	0.18	0.2
CTSM13.24	13 12 28.60	–25 17 00.7	19025	0.063460	17.5V	0.18	—
FAIRALL0107*	21 35 29.50	–62 30 07.2	18275	0.060959	16.7b	0.44	0.34
HE0348–5353	03 49 28.00	–53 44 47.0	>30000	0.130000	15.7V	—	—
HE1107+0129*	11 10 12.07	+01 13 27.8	28630	0.095500	16.5g	0.31	0.30
IRAS04576+0912*	05 00 20.77	+09 16 55.6	10822	0.036098	16.6V	—	—
IRAS16355–2049*	16 38 30.92	–20 55 24.6	7906	0.026372	14.5V	0.30	—
IRAS20520–2329	20 54 57.30	–23 18 24.0	>30000	0.206300	16.3R	—	—
MCG–04.24.017*	10 05 55.37	–23 03 25.0	3842	0.012816	14.5V	1.0	0.5
RBS0219	01 35 27.00	–04 26 35.0	>30000	0.154700	16.5V	0.23	0.20
RHS56	20 39 27.19	–30 18 52.2	23709	0.079085	16.0V	—	—
RXJ0024.7+0820*	00 24 45.70	+08 20 56.9	20086	0.067000	18.2V	—	—
RXJ0323.2–4931*	03 23 15.35	–49 31 06.7	21285	0.071000	12.4R	0.48	0.30
RXJ0902.5–0700*	09 02 33.57	–07 00 04.3	26715	0.089112	17.7b	—	—
RXJ2301.8–5508*	23 01 52.01	–55 08 31.1	>30000	0.141000	14.7R	—	—
SDSSJ134524.70–025939.8*	13 45 24.70	–02 59 39.8	25566	0.085279	16.8g	0.27	0.17
SDSSJ161227.83+010159.8	16 12 27.84	+01 01 59.9	29109	0.097096	18.6g	—	—
SDSSJ225452.22+004631.4*	22 54 52.22	+00 46 31.4	27202	0.090735	17.3g	—	—
V961349–439*	13 52 59.63	–44 13 25.4	15589	0.052000	15.4V	—	—
WPV85007*	00 39 15.85	–51 17 01.5	8577	0.028610	15.8V	0.18	0.15
Zw049.106*	15 17 51.70	+05 06 27.7	11626	0.038780	15.6V	0.70	0.43

Table A.2: Observed galaxies. Column 1: Galaxy name. Column 2: Observation date (mm/dd/yyyy) of the blue range spectra ($\sim 4300 \text{ \AA} - \sim 5200 \text{ \AA}$). Column 3: Exposure time of the blue range. Column 4: Observation date (mm/dd/yyyy) of the red range spectra ($\sim 5800 \text{ \AA} - \sim 6800 \text{ \AA}$). Column 5: Exposure time of the red range. Column 6: S/N in the range $4750 \text{ \AA} - 4800 \text{ \AA}$. The red spectra of galaxies labeled with "*" are available in Schmidt et al. (2016). The S/N of the object labeled with "†" was measured in the spectral range $5070 \text{ \AA} - 5120 \text{ \AA}$.

Galaxy name	Date blue range	Exposure time (s)	Date red range	Exposure time (s)	S/N (4750 \AA – 4800 \AA)
1RXS J040443.5–295316*	10/26/2014	3×2400	—	—	10
2MASX J01413249–1528016*	10/28/2014	3×2400	—	—	12
2MASXJ05014863–2253232*	10/13/2015	3×2400	—	—	21
2MASXJ08173955–0733089	04/12/2013	3×2400	04/10/2013	3×1800	7
2MASX J21531910–1514111*	10/26/2014	2×2400	—	—	11
2MASXJ21565663–1139314*	10/16/2015	3×2400	—	—	21
CTSH34.06	10/19/2015	4×2400	—	—	22
CTSM02.47*	04/26/2012	3×1800	—	—	6
CTSM13.24	04/29/2014	3×2400	04/10/2013	3×1800	13
FAIRALL0107*	10/13/2015	4×1800	—	—	24
HE0348–5353	10/28/2014	3×2400	10/27/2014	3×2400	11
HE1107+0129*	04/08/2013	3×1800	—	—	17
IRAS04576+0912*	10/18/2015	3×2400	—	—	21
IRAS16355–2049*	04/23/2012	3×1800	—	—	20
IRAS20520–2329†	05/02/2014	2×1800	—	—	15
MCG–04.24.017*	04/25/2012	3×1800	—	—	21
RBS0219	10/17/2015	3×2400	08/18/2012	3×1800	31
RHS56	10/17/2015	4×2400	08/18/2012	2×1800	25
RX J0024.7+0820*	10/26/2014	3×2400	—	—	10
RX J0323.2–4931*	10/17/2015	4×2400	—	—	17
RX J0902.5–0700*	04/08/2013	3×1800	—	—	8
RX J2301.8–5508*	10/28/2014	3×2400	—	—	23
SDSS J134524.69–025939.8*	04/08/2013	2×2400	—	—	11
SDSSJ161227.83+010159.8	07/13/2015	3×1800	07/12/2015	4×1800	5
SDSS J225452.22+004631.4*	09/10/2013	2×2400	—	—	10
V961349–439*	04/07/2013	3×1800	—	—	19
WPV8507*	10/18/2015	3×2400	—	—	28
Zw049.106*	04/29/2014	3×2700	—	—	13

Table A.3: FWHM of the BC of $H\alpha$, luminosity of that component, and black hole masses. Parameters of galaxies labeled with "*" correspond to $BC\beta$, with the black hole estimation through this line.

Galaxy	FWHM _{BCα} (km s ^{−1})	log L _{BCα} (erg s ^{−1})	log(M _{BH}) (M _⊙)
2MASXJ08173955–0733089	2697	41.12	6.7
CTSH34.06*	1685	40.23	5.8
CTSM13.24	2715	41.05	6.7
HE0348–5353	3065	42.03	7.3
IRAS20520–2329*	2016	41.82	6.8
RBS0219	2721	41.94	7.2
RHS56	4159	41.68	7.4
SDSSJ161227.83+010159.8	1198	40.66	5.7

Table A.4: Parameters of the [OIII] λ 5007 emission profile. Column 1: Galaxy name. Columns 2 and 3: FWHM and flux of the core component of [OIII] λ 5007. Columns 4 and 5: FWHM and flux of the asymmetric component of [OIII] λ 5007. Column 6: Radial velocity difference between the centroid of the narrow component of H β and the centroid of the core component of [OIII] λ 5007. Column 7: Radial velocity difference between the centroid of the asymmetric component and the core component of [OIII] λ 5007. Column 8: Radial velocity difference between the blue end of the blue wing and the centroid of the core component of [OIII] λ 5007 (equation 2). Column 9: Logarithm of the Eddington ratio. All presented FWHM were corrected by instrumental broadening and are in units of km s $^{-1}$, similar to the three presented radial velocity differences. All fluxes are in units of 10^{-15} erg cm $^{-2}$ s $^{-1}$.

Galaxy name	FWHM _{[OIII]λ5007} (km s $^{-1}$)	Flux _{[OIII]λ5007} (10^{-15} erg cm $^{-2}$ s $^{-1}$)	FWHM _{[OIII]λ5007} (km s $^{-1}$)	Flux _{[OIII]λ5007} (10^{-15} erg cm $^{-2}$ s $^{-1}$)	Δv (km s $^{-1}$)	Δv_{asym} (km s $^{-1}$)	total asymmetry (km s $^{-1}$)	L_{bol}/L_{Edd}
IRXSJ040443.5-295316	124	2.70	314	6.88	88	-173	-486	0.01
2MASXJ01413249-1528016	260	2.03	597	1.22	-58	-368	-965	-0.29
2MASXJ05014863-2253232	262	1.35	761	12.72	45	-201	-963	-0.20
2MASXJ08173955-0733089	261	1.18	1006	0.92	222	-123	-1129	-0.93
2MASXJ21531910-1514111	415	1.57	933	1.01	-113	-180	-1113	-0.18
2MASXJ21565663-1139314	290	19.04	795	5.12	10	-62	-857	0.18
CTSH34.06	283	18.75	571	7.86	27	-121	-692	0.39
CTSM02.47	312	4.13	369	1.07	17	-308	-677	-0.54
CTSM13.24	306	4.15	789	5.81	180	-153	-943	-0.35
FAIRALL0107	364	2.35	971	3.65	5	-35	-1005	-0.46
HE0348-5353	365	7.33	930	12.70	-3	-165	-1095	-0.06
HE1107+0129	270	4.40	846	2.47	9	-175	-1020	-0.45
IRAS04576+0912	373	7.98	1184	9.78	-220	-498	-1682	-0.37
IRAS16355-2049	341	67.50	940	101.40	-36	-224	-1164	-0.40
IRAS20520-2329	319	2.51	1010	3.93	-59	-267	-1277	0.36
MCG-04.24.017	203	17.85	621	13.81	7	-25	-646	-0.60
RBS0219	580	5.19	954	5.87	-1	-566	-1520	-0.02
RHS56	255	11.27	1114	20.09	23	-118	-1232	-0.38
RXJ0024.7+0820	365	2.22	370	0.27	-43	-711	-1081	-0.33
RXJ0323.2-4931	218	3.38	1041	1.69	176	-144	-1185	-0.98
RXJ0902.5-0700	324	4.49	1358	2.75	-68	-332	-1689	-0.54
RXJ2301.8-5508	476	6.21	608	1.24	-105	-893	-1501	-0.59
SDSSJ134524.70-025939.8	336	4.23	900	2.92	-104	-199	-1099	-0.27
SDSSJ161227.83+010159.8	236	0.63	732	0.93	-38	-39	-771	0.16
SDSSJ225452.22+004631.4	135	0.51	1431	1.96	288	-109	-1540	-0.83
V961349-439	359	5.74	684	6.59	-405	-126	-810	-0.63
WPV85007	315	8.72	1041	6.21	-186	-432	-1473	-0.78
Zw049.106	369	7.53	905	3.19	46	-302	-1207	-0.38

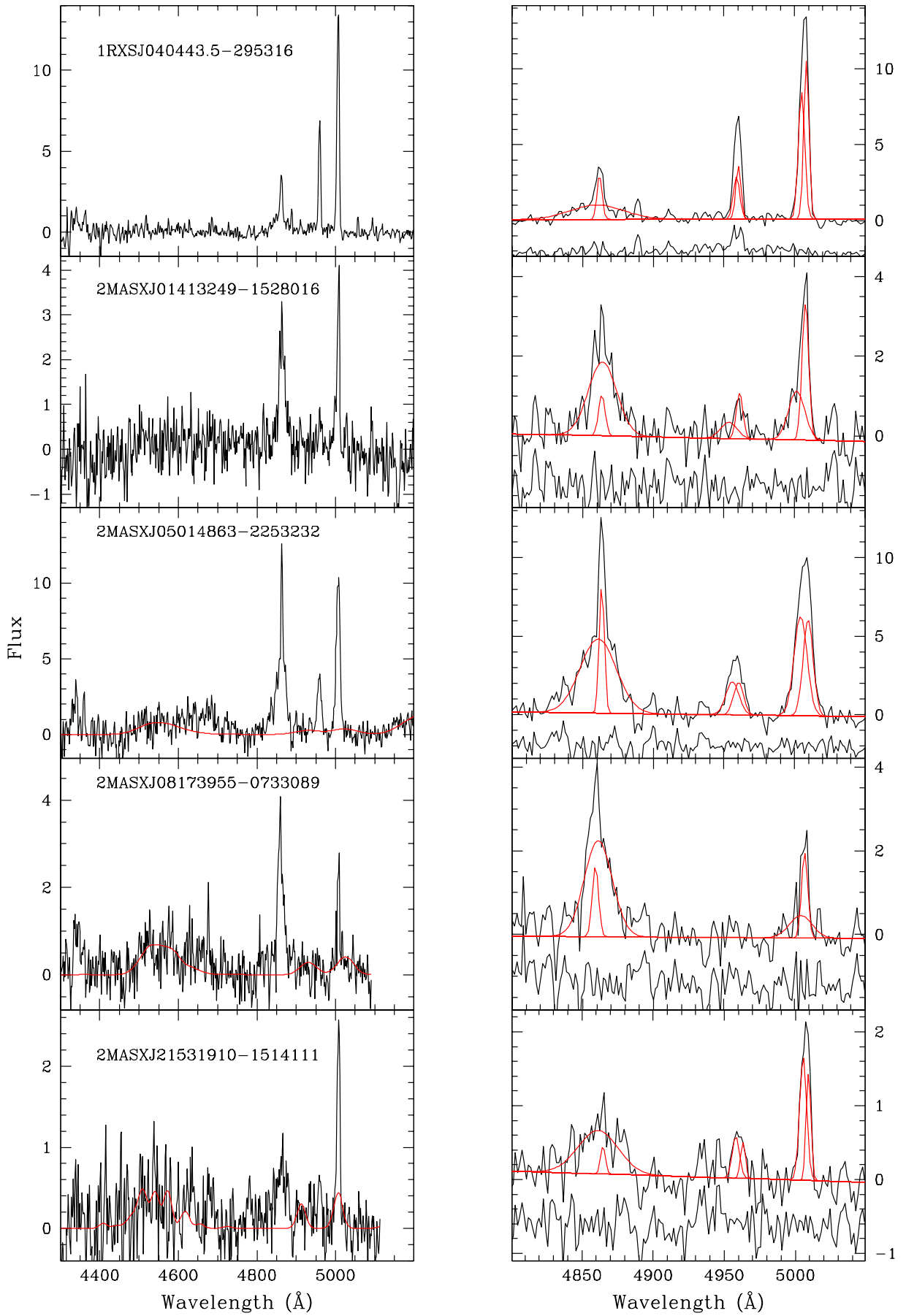


Fig. A.1: Left: Continuum subtracted spectra of the NLS1 in the range 4300 \AA - 5200 \AA . Overlaid in red is the Fe multiplets fitting. Right: Gaussian decomposition of the H β and [OIII] $\lambda\lambda 4959, 5007$ emission lines of the Fe subtracted spectra. The individual components are shown in red lines and the residuals are plotted at the bottom of each panel. Fluxes are in arbitrary units.

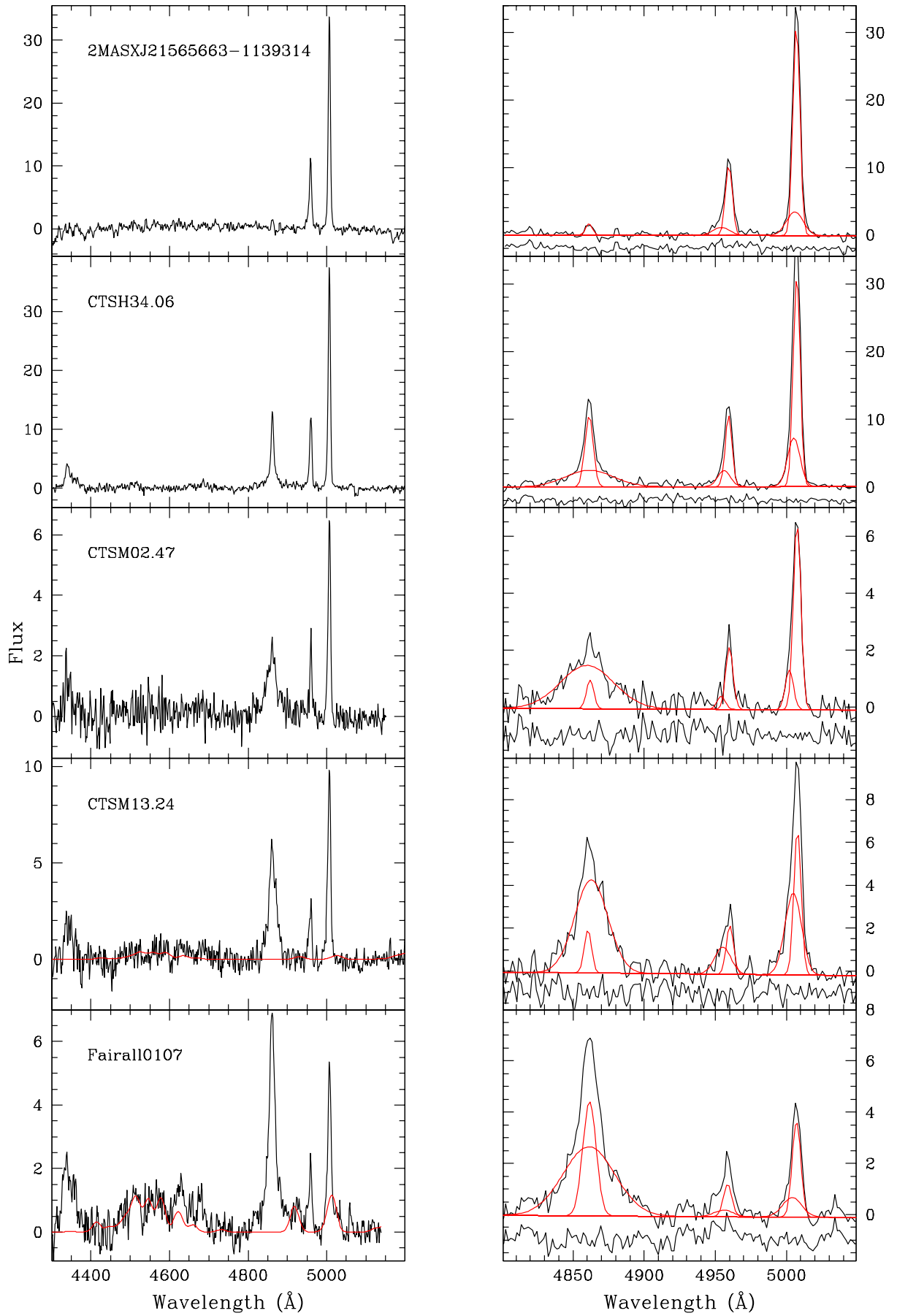


Fig. A.1: Continued

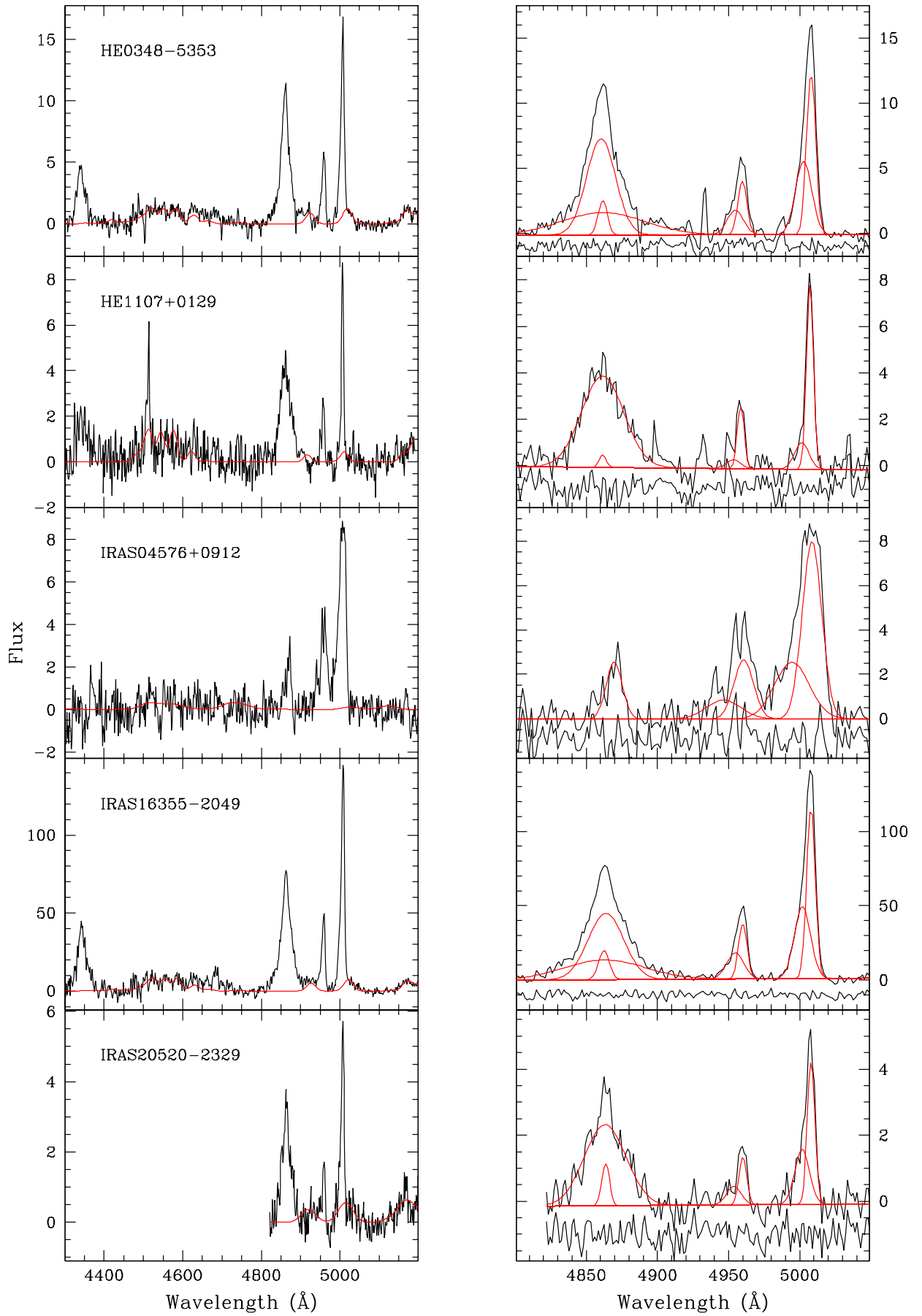


Fig. A.1: Continued

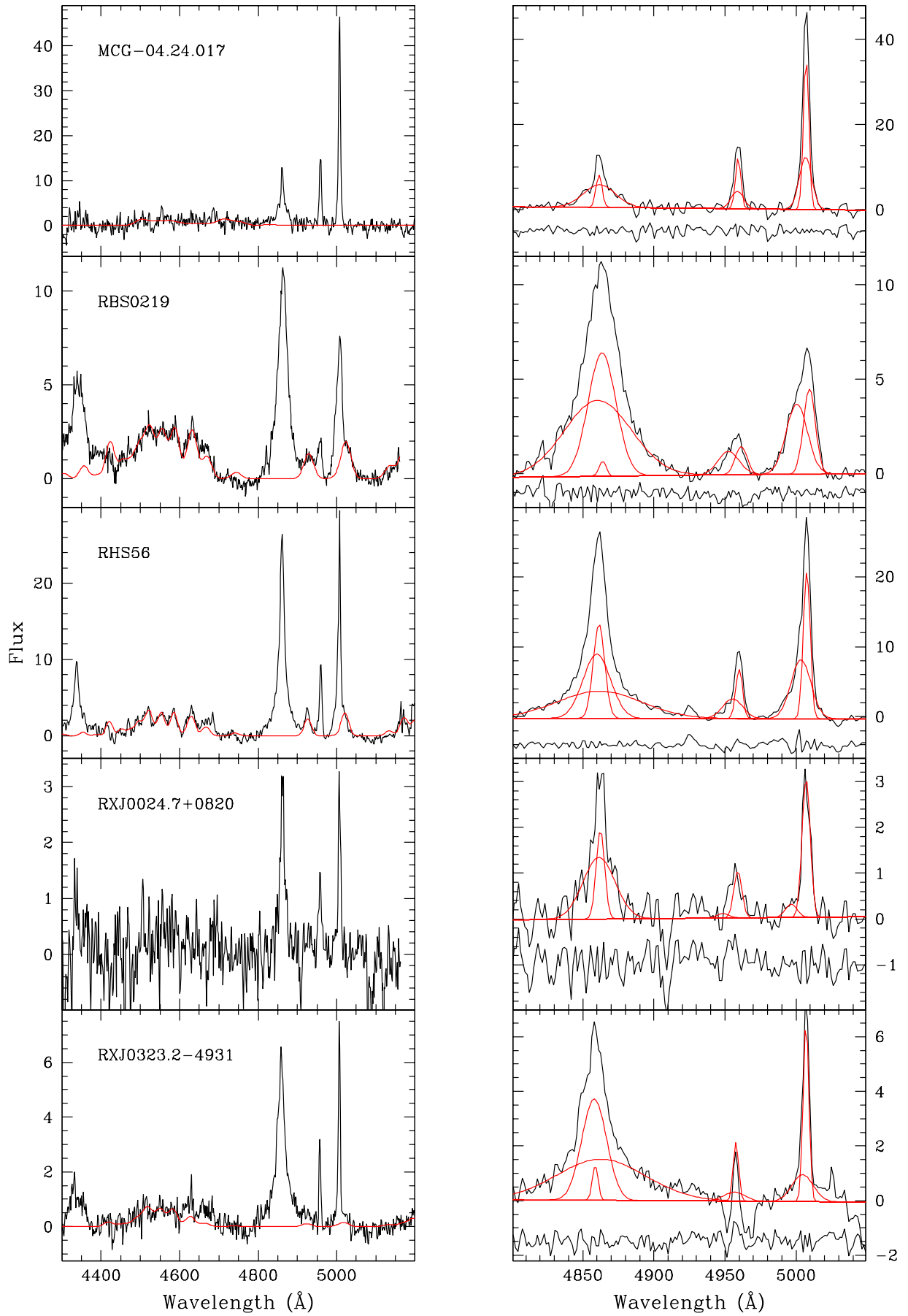


Fig. A.1: Continued

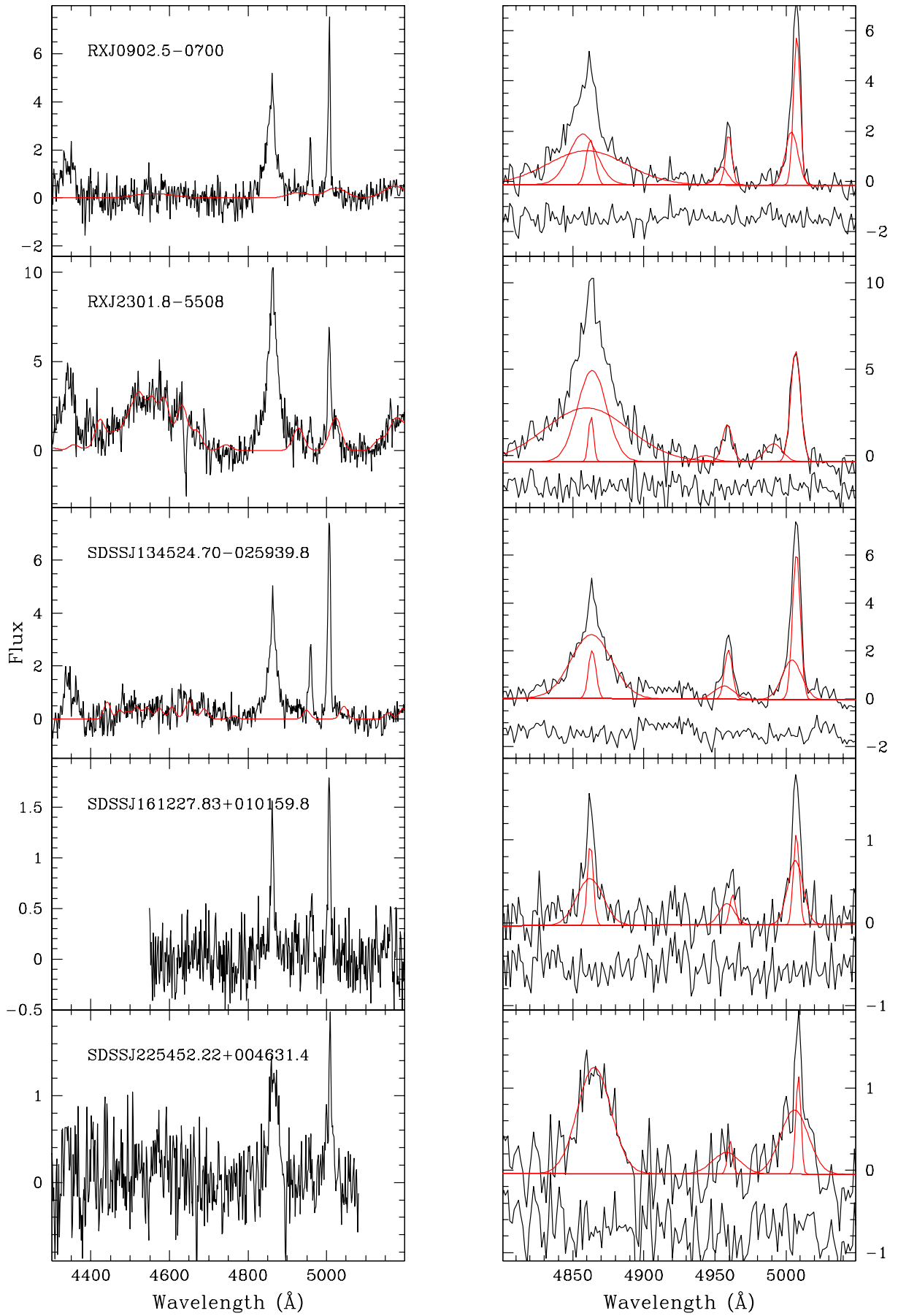


Fig. A.1: Continued

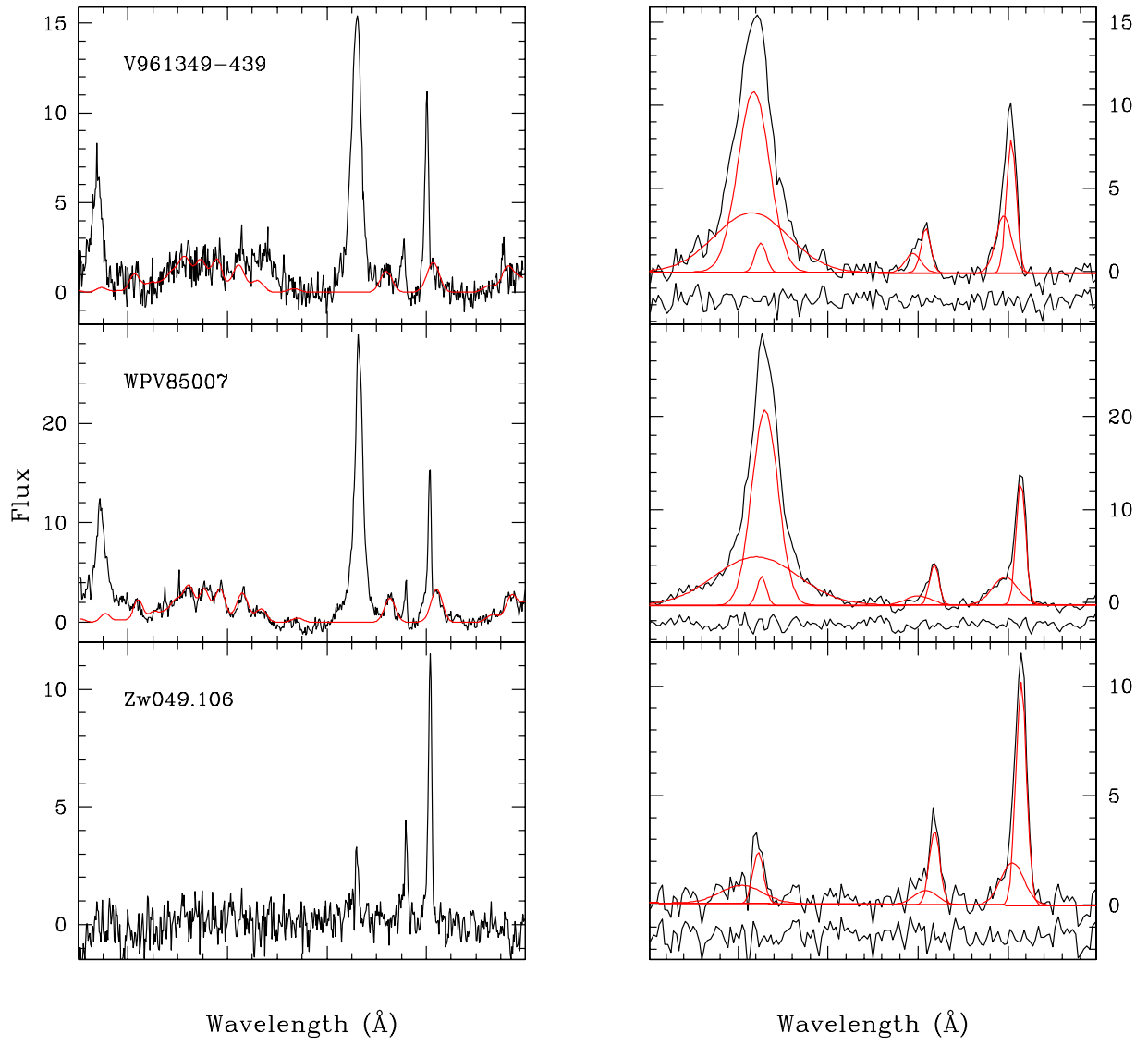


Fig. A.1: Continued

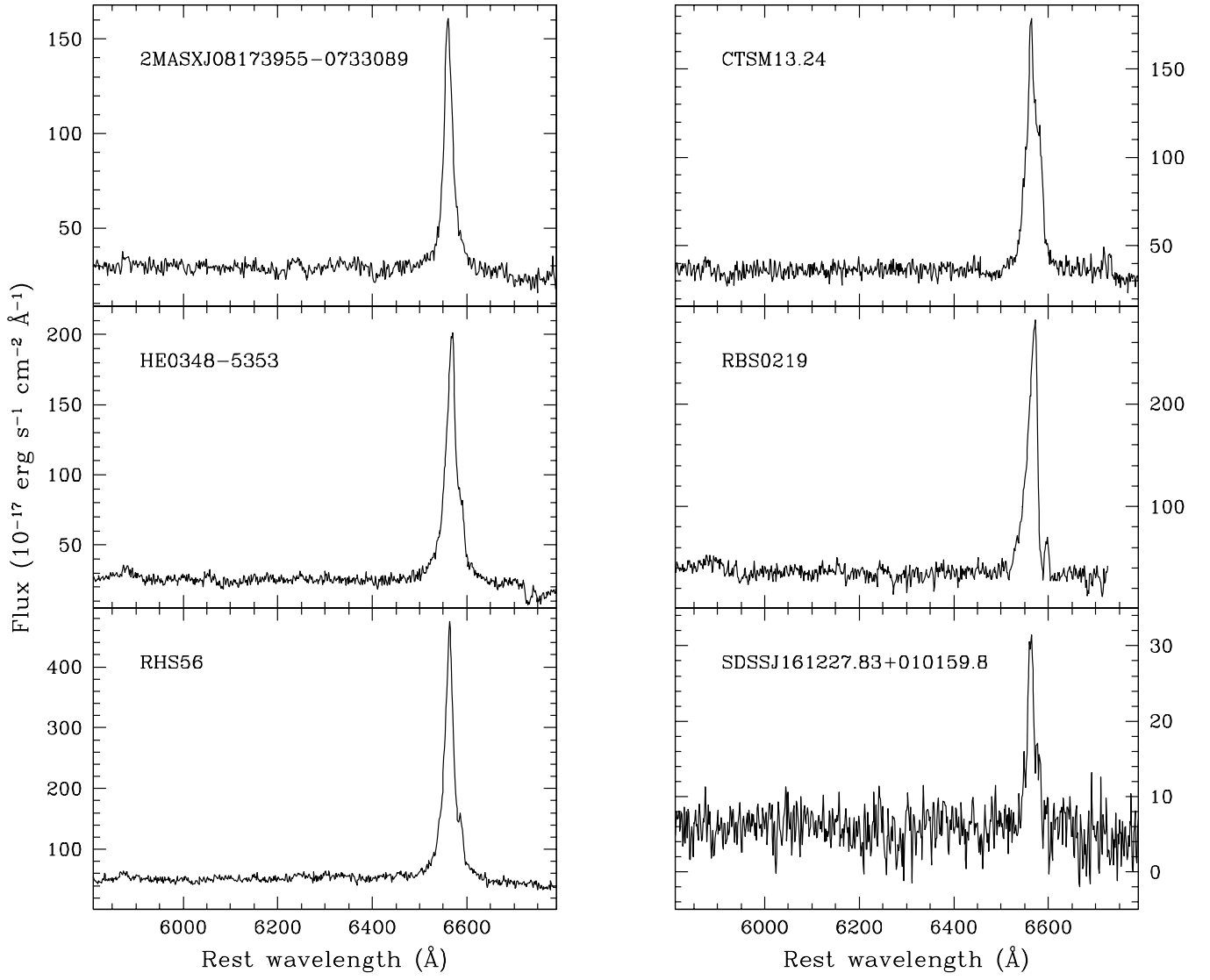


Fig. A.2: Observed spectra of the NLS1 in the range 5800Å - 6800Å. Fluxes are in units of $10^{-17} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Å}^{-1}$.